

Multi-band study of RX J0838–2827 and XMM J083850.4–282759: a new asynchronous magnetic cataclysmic variable and a candidate transitional millisecond pulsar

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ABSTRACT

In search for the counterpart to the *Fermi*-LAT source 3FGLJ0838.8–2829, we performed a multi-wavelength campaign, in the X-ray band with *Swift* and *XMM-Newton*, performed infrared, optical (with OAGH, ESO-NTT and IAC80) and radio (ATCA) observations, as well as analysed archival hard X-ray data taken by *INTEGRAL*. We report on three X-ray sources consistent with the position of the *Fermi*-LAT source. We confirm the identification of the brightest object, RX J0838–2827, as a magnetic cataclysmic variable, that we recognize as an asynchronous system (not associated with the *Fermi*-LAT source). RX J0838–2827 is extremely variable in the X-ray and optical bands, and timing analysis reveals the presence of several periodicities modulating its X-ray and optical emission. The most evident modulations are interpreted as due to the binary system orbital period of ~ 1.64 h and the white dwarf spin period of ~ 1.47 h. A strong flux modulation at ~ 15 h is observed at all energy bands, consistent with the beat frequency between spin and orbital periods. Optical spectra show prominent H β , HeI and HeII emission lines Doppler-modulated at the orbital period and at the beat period. Therefore, RX J0838–2827 accretes through a diskless configuration and could be either a strongly asynchronous polar or a rare example of a pre-polar system in its way to reach synchronism. Among the other two X-ray sources, XMM J083850.4–282759 showed a variable X-ray emission, with a powerful flare lasting ~ 600 s, similar to what is observed in transitional millisecond pulsars during the sub-luminous disc state: that would possibly associate this source with the *Fermi*-LAT source.

Key words: accretion, accretion discs – novae, cataclysmic variables – pulsars: general – white dwarfs – X-rays: individual: RX J0838–2827– X-rays: individual: XMM J083850.4–282759

1 INTRODUCTION

1.1 Cataclysmic variables

Cataclysmic variables (hereafter CVs) are close binary systems hosting a white dwarf (WD) and a low-mass main sequence star typically filling its Roche lobe and transferring material toward the compact object (see Warner 2003 for a review). About 20–25 per cent of the known CVs harbour WDs with magnetic fields in the 10^5 – 10^8 G range, and are known as magnetic CVs (mCVs; Ferrario, de Martino & Gänsicke 2015). The mCVs can be divided into two main groups: the intermediate polars (IPs) and the polars. The former are characterised by asynchronously rotating WDs ($P_{\text{spin}} \ll P_{\text{orb}}$) with spin periods ranging from few hundreds to a thousands of seconds, and orbital periods from one to tens of hours; typically, they accrete via a truncated accretion disc. Due to the lack of synchronisation and of optical/nIR circular polarisation, the WD is believed to possess moderately low magnetic fields ($\leq 10^7$ G). The Polars instead consist of orbitally-locked rotating WDs ($P_{\text{spin}} = P_{\text{orb}} \sim \text{hrs}$) and are circularly polarised in the optical/nIR (see Cropper 1990). Their magnetic field as measured from cyclotron emission ranges between ~ 10 – 300 MG, preventing the formation of an accretion disc; thus, accretion proceeds directly from the donor star onto the WD magnetic poles through a stream.

A handful of polars were found to be slightly desynchronised ($\leq 2\%$), namely V1500 Cyg, BY Cam, V1432 Aql and CD Ind (Campbell & Schwope 1999). Another candidate, SWIFT J2319.4+2619, was found recently (Shafer et al. 2008). Since V1500 Cyg exploded as a nova in 1975, these asynchronous polars are believed to have undergone a nova explosion in the past and they will return to synchronism in hundreds of years (Boyd et al. 2014; Harrison & Campbell 2016). However, only V1500 Cyg is known as a nova, and searches for nova shells resulted in no detections (Pagnotta & Zurek 2016).

Recently, two systems were found to be desynchronised at a much higher level (i.e. $\gg 2\%$), namely Paloma (Schwarz et al. 2007) and IGR J19552+0044 (Bernardini et al. 2013). These systems might be either regarded as IPs with a very low degree of asynchronism or polars with a large degree of asynchronism. In the former case these systems are not expected to reach synchronism, while in the latter case they may represent the true polar progenitors (see Norton et al. 2004). Despite their small number, asynchronous polars represent excellent laboratories to study the physics of mass transfer in the case of a varying magnetic field geometry, especially because of the continuous changes of the orientation between the WD magnetic field axis and the accretion flow.

1.2 Transitional millisecond pulsars

The short spin periods of binary millisecond pulsars are the outcome of the accretion onto the neutron star (NS) of the mass transferred by a low mass late type companion star, through an accretion disc (Alpar et al. 1982). After a Gyr-long mass accretion phase during which the binary system shines as a bright low mass X-ray binary (NS-LMXB), the mass transfer rate declines and allows the activation of a—now rapidly spinning—radio/gamma-ray millisecond pulsar (MSP), powered by the rotation of its magnetic field. The tight link existing between radio MSPs and NS-LMXBs has been recently demonstrated by the discovery of three transitional millisecond pulsars (PSR J1023+0038, Archibald et al. 2009; IGR

J18245–2452 in the globular cluster M28, Papitto et al. 2013; XSS J12270–4859, de Martino et al. 2014, Bassa et al. 2014).

These sources have been observed to switch between accretion and rotation-powered emission on timescales ranging from a couple of weeks to months, showing that such state transitions may take place on time scales compatible with those of the variations of the mass accretion rate onto the NS. At high mass inflow rates, the radio pulsar is shut-off and the system is bright in the X-rays ($L_X > 10^{36}$ erg s $^{-1}$). At low mass inflow rates, the magnetosphere expands up to the light cylinder activating the radio pulsar, the disc disappears and the system is faint in the X-rays ($L_X \sim 10^{32}$ erg s $^{-1}$).

Surprisingly, in addition to the X-ray outbursts accreting state and the radio pulsar state, the three known transitional MSPs have been observed into another extremely peculiar *sub-luminous* disc state with $L_X \sim 10^{33}$ erg s $^{-1}$ (see also Linares 2014). During these states, both PSR J1023+0038 (Archibald et al. 2015) and XSS J1227–4859 (Papitto et al. 2015) showed that part of the disc material is accreted onto the NS surface. These sub-luminous states are accompanied by X-ray flaring activity, a sizeable gamma-ray emission and a flat-radio spectrum that are typical jet signatures in accreting compact objects, suggesting that large mass outflows could be launched by the fast rotating propelling magnetosphere of these pulsars (Papitto et al. 2014; Papitto & Torres 2015; but see also Takata et al. (2014) for an alternative modelling based on a pulsar wind/disk shock).

1.3 The ROSAT source: 1RXS J083842.1–282723

1RXS J083842.1–282723 (RX J0838–2827 hereafter) was discovered by ROSAT, and tentatively classified as a CV by Masetti et al. (2013) based on Balmer and helium emission lines observed in its optical spectrum. Since two of the three transitional MSPs were formerly misclassified as CVs because of their CV-like optical spectra, searches within the error circles of newly discovered but still unidentified *Fermi*-LAT sources are being carried out despite early tentative classifications based on optical spectroscopy. In particular, the spatial coincidence with the unidentified gamma-ray source 3FGL J0838.8–2829 in the *Fermi* Large Area Telescope (LAT) 4-year point source catalog (Acero et al. 2015) and the variability of the X-ray emission detected on a time scale of a few hours in archival *Swift* observations, made the nature of the source rather controversial.

We report in this paper on a large unprecedented multi-band campaign of the field of 3FGL J0838.8–2829 in the radio, infrared, optical and X-ray bands. In Sections 2, 3, and 4, we report on the details of the observations of the field of 3FGL J0838.8–2829. In Section 5, data analysis and results on RX J0838–2827 are collected. In Section 6, we report on the analysis of the other X-ray sources associated with 3FGL J0838.8–2829 and in particular on the flaring source XMM J083850.4–282759, which might be a new transitional millisecond pulsar possibly associated with the gamma-ray emission. Discussion follows in Section 7.

2 X-RAY OBSERVATIONS

We report here and in the following few sections on the multi-band data analysis and results of RX J0838–2827, while in Section 6 we report on the other X-ray sources in the field, in particular on the flaring X-ray source XMM J083850.4–282759.

2.1 XMM–Newton

Two XMM–Newton observations were carried out on 2015 October 20–21 (ID: 0764420101; PI: Rea) and December 2–3 (ID: 0790180101; thanks to XMM Director’s Time), respectively, and in both cases RX J0838–2827 was placed at the aim point of the European Photon Imaging Cameras (EPIC). The pn camera (Strüder et al. 2001) was configured in large window mode (LW; 47.7-ms time resolution), whereas the two MOS cameras (Turner et al. 2001) were operated in small window mode (SW; 0.3-s time resolution). The thin optical blocking filter was positioned in front of the cameras in both observations. The Reflection Grating Spectrometers (RGS; den Herder et al. 2001) were operated in standard spectroscopy mode. The journal of the two observations is reported in Table 1.

2.1.1 EPIC data

We processed the raw observation data files using the EPPROC (for pn data) and EMPROC (for MOS data) tasks of the XMM–Newton Science Analysis System (SAS¹, version 15.0), with the calibration files in the database released in 2016 January (XMM-CCF-REL-332).

The inspection of the EPIC light curves binned at 10 s revealed episodes of strong flaring particle background in both observations. For the spatial and spectral analysis, we discarded the intervals affected by flaring background, which were pinpointed by applying intensity filters on the light curves, following the procedure described by De Luca & Molendi (2004). This reduced the effective exposure time to approximately 38.9 (pn), 48.5 (MOS 1) and 48.4 (MOS 2) ks for the first observation, and to 64.6 (pn), 71.5 (MOS 1) and 71.2 (MOS 2) ks for the second observation.

We ran the SAS source detection script EDETECT_CHAIN on the combined event lists from the two observations. This returned for RX J0838–2827 the position RA = 08^h38^m43^s.20, Dec = –28°27′01″.44 (J2000.0), with an uncertainty of 1.4 arcsec (radius, at the 90 per cent confidence level).² This is fully consistent with the more accurate position of the optical counterpart, RA = 08^h38^m43^s.37, Dec = –28°27′01″.5, with an uncertainty of about 0.2 arcsec (Masetti et al. 2013). The 0.2–10 keV image of the field of 3FGL J0838.8–2829 is shown in Fig. 1, and RX J0838–2827’s light-curves in Fig. 2. It was created by stacking the images of the EPIC cameras cleaned from bad pixels and hot columns from both observations. Apart from RX J0838–2827, two other X-ray sources were identified in the 68 per cent error circle of 3FGL J0838.8–2829. In the following we will dub them as XMM J083850.4–282759 and XMM J083842.7–283832 (see Table 4 for details on their position and their X-ray properties).

For the study of RX J0838–2827, we collected the source counts within a circle with a 40 arcsec radius for the pn data, and 30 arcsec for the MOS data. The background was extracted from circles located on the same CCD as the source (so as to guarantee similar low-energy noise subtraction) and avoiding detector areas contaminated by possible out-of-time events from the source (in the pn) or too near to the CCD edges. On the other hand, data products for the other two sources were extracted only from the event

lists of the pn, because both fell outside the central window of the MOS instruments.

We restricted the analysis to photons in the 0.2–10 keV energy range with PATTERN ≤ 4 for the pn and ≤ 12 for the MOS, and FLAG = 0 (we also checked that the data were not affected by pile-up). For the spectral analysis of RX J0838–2827 we used only the data acquired by the pn camera, which provides the spectra with the highest counting statistics, to avoid systematic errors introduced by cross-calibration uncertainties (we checked that the MOS spectra are in agreement with the results from the pn camera). We generated the spectral redistribution matrices and ancillary response files for each of the three sources with the RMFGEN and ARFGEN tools, respectively. The fits were performed with the XSPEC package (version 12.9.0; Arnaud 1996), using the χ^2 statistics. For the timing analysis, we converted the photon arrival times to Solar system barycentre reference frame using the SAS task BARYCEN and the DE-200 Solar System ephemeris.

2.1.2 RGS data

We processed the data using the task RGSPROC of SAS to produce calibrated and concatenated photon event lists, spectra and response matrices. We focused our analysis on the first order spectra alone, which are the best calibrated and have a higher number of counts. First, we subtracted the model background spectra, which were created by the standard RGS pipeline and are template background files based on the count rate in CCD 9. We then combined the spectra using RGS COMBINE, which appropriately accounts for the response matrices and backgrounds of the different spectrometers, and rebinned the resulting spectra to have at least one count per energy bin. The analysis was limited to the 0.35–2.5 keV energy interval (5–38 Å), where the calibration of the response is most accurate (see Fig. 3).

2.2 Swift X-Ray Telescope

The X-ray Telescope (XRT; Burrows et al. 2005) on board the Swift satellite observed the field of RX J0838–2827 for a total of 14 times between 2010 September 8 and 2016 January 20 (see Table 1 for a log of the observations, and Fig. 4). The first 7 observations were carried out in photon counting (PC) mode, which provides a two-dimensional image and a time resolution of 2.5073 s. The last 7 observations were instead performed in windowed timing (WT) mode, i.e., data were transferred from the imaging array to the frame store array with an integration time of 1.78 ms, at the expense of one dimension of spatial resolution (see Hill et al. 2004 for a detailed description of the XRT readout modes).

The data were analysed primarily to study the X-ray long-term variability of RX J0838–2827. We processed the data with standard screening criteria and generated exposure maps with the task XRTPIPELINE (version 0.13.2) from the FTOOLS package (Blackburn 1995), using the optical position of the source (Masetti et al. 2013) and the spacecraft attitude file. We selected events with grades 0–12 and 0 for the PC and WT data, respectively, and extracted the source and background spectra using XSELECT (v. 2.4). We accumulated the source counts from a circular region with a radius of 20 pixels (one XRT pixel corresponds to about 2.36 arcsec). To estimate the background in the PC-mode data, we extracted the events within an annulus centred on the source position with inner and outer radius of 40 and 80 pixels, respectively. For the WT-mode data we opted for a circle far from the target and of the same size as the source one, owing to the small extent of the WT window.

¹ See <http://xmm.esac.esa.int/sas/>.

² The uncertainty on the position is evaluated as the quadratic sum of the statistical uncertainty given by the task EDETECT_CHAIN and the systematic absolute pointing uncertainty of XMM–Newton.

Table 1. Log of the X-ray observations.

Satellite	Obs. ID	Date (YY/MM/DD)	Start – End time of exposure (MJD)	Instrument	Mode	Exposure ^a (ks)	RX J0838–2827 count rate ^b (counts s ⁻¹)
<i>XMM-Newton</i>	0764420101	2015/10/20-21	57 315.4250 – 57 316.0116	EPIC pn	LW	47.2	2.053 ± 0.007
				EPIC MOS 1	SW	51.1	0.5957 ± 0.0004
				EPIC MOS 2	SW	51.0	0.5743 ± 0.0003
				RGS 1	spectroscopy	52.7	0.0447 ± 0.0007
				RGS 2	spectroscopy	52.8	0.0456 ± 0.0008
<i>XMM-Newton</i>	0790180101	2015/12/02-03	57 358.9411 – 57 359.8111	EPIC pn	LW	69.1	1.714 ± 0.007
				EPIC MOS 1	SW	74.0	0.529 ± 0.003
				EPIC MOS 2	SW	74.0	0.533 ± 0.003
				RGS 1	spectroscopy	76.3	0.0422 ± 0.0005
				RGS 2	spectroscopy	76.4	0.0429 ± 0.0007
<i>Swift</i>	00041343002	2010/10/02	55 471.7469 – 55 471.9604	XRT	PC	4.1	0.086 ± 0.005
	00046948001	2011/12/18	55 913.4285 – 55 913.5048	XRT	PC	2.0	0.074 ± 0.006
	00046948002	2014/10/07	56 937.9180 – 56 937.9999	XRT	PC	3.0	0.069 ± 0.005
	00046948003	2014/10/11	56 941.3212 – 56 941.8555	XRT	PC	2.6	0.061 ± 0.005
	00046948004	2014/10/15	56 945.1198 – 56 945.5923	XRT	PC	2.7	0.082 ± 0.006
	00046948005	2014/10/19	56 949.2479 – 56 949.7930	XRT	PC	2.5	0.127 ± 0.007
	00041343003	2015/10/14	57 309.4321 – 57 309.5680	XRT	PC	2.0	0.18 ± 0.01
	00041343004	2015/12/27-28	57 383.0380 – 57 384.1743	XRT	WT	10.4	0.217 ± 0.005
	00041343005	2015/12/31	57 387.4330 – 57 387.6291	XRT	WT	1.4	0.37 ± 0.02
	00046948006	2016/01/12	57 399.0082 – 57 399.7305	XRT	WT	4.0	0.312 ± 0.009
	00046948007	2016/01/13	57 400.0037 – 57 400.5470	XRT	WT	0.5	0.13 ± 0.02
	00046948008	2016/01/14	57 401.0678 – 57 402.0000	XRT	WT	3.5	0.151 ± 0.007
	00046948009	2016/01/17	57 404.3239 – 57 404.6413	XRT	WT	2.8	0.157 ± 0.008
	00046948010	2016/01/20	57 407.3096 – 57 407.8395	XRT	WT	3.7	0.190 ± 0.008

^a Dead-time corrected on-source exposure.

^b Average background-subtracted count rate in the 0.2–10 keV energy band for the EPIC instruments, in the 0.35–2.5 keV energy band for the combined RGS first order data sets, and in the 0.3–10 keV energy band for the *Swift* XRT data.

We created the observation-specific ancillary response files (using exposure maps) with XRTMKARF (v. 0.6.3), which corrects for the loss of counts due to hot columns and bad pixels, and accounts for different extraction regions, telescope vignetting and PSF corrections. We then assigned the appropriate redistribution matrix available in the HEASARC calibration data base,³ excluded bad spectral channels, and grouped the background-subtracted spectra to have at least 20 counts in each spectral bin. We limited the spectral analysis to the 0.3–10 keV energy range for PC-mode data and to the 0.7–10 keV interval for the WT-mode data, owing to known calibration issues in WT that result in spurious bumps and/or turn-ups in the spectra at low energy.⁴

2.3 INTEGRAL

We analyzed all publicly available *INTEGRAL* data collected in the direction of the source in the past 13 yrs of the mission. We considered all *INTEGRAL* science windows (ScW), i.e. the different pointings lasting each ~ 2 –3 ks, in which the source was observed within 12 degrees off axis from the satellite aim point for the IBIS/ISGRI instrument (Ubertini et al. 2003; Lebrun et al. 2003) and 3.5 degrees off axis for the two JEM-X instruments (Lund et al. 2003). All data were processed and analyzed by making use of the Off-line Scientific Analysis software (OSA) version 10.2 distributed by the ISDC (Courvoisier et al. 2003).

³ See <http://www.swift.ac.uk/analysis/xrt/rmfarf.php>.

⁴ See http://www.swift.ac.uk/analysis/xrt/digest_cal.php.

We found that the region around 3FGLJ0838.8–2829 was probed by *INTEGRAL* for a total of 892 ks effective exposure with IBIS/ISGRI, and 15.5 ks (10.0 ks) with JEM-X1 (JEM-X2). The dataset spanned from the satellite revolution 28 (starting on 52644 MJD) to 1625 (ending on 57381 MJD). We built the IBIS/ISGRI mosaics in the two 20–40 keV and 40–80 keV energy bands, as well as the JEM-X mosaics in the 3–10 keV energy band. RX J0838–2827 was not detected by these instruments. We estimated with the OSA `mosaic_spec` tool a 3σ upper limit on the source hard X-ray flux of 1 mCrab (roughly 8×10^{-12} erg cm⁻² s⁻¹) in the 20–40 keV energy band and 1.5 mCrab (roughly 10^{-11} erg cm⁻² s⁻¹) in the 40–80 keV energy band. The corresponding upper limit in the 3–10 keV energy band was of 5 mCrab (roughly 7×10^{-11} erg cm⁻² s⁻¹), consistent with the *XMM-Newton* findings.

3 INFRARED AND OPTICAL OBSERVATIONS

3.1 OAGH

Near-IR observations of the field of the 3FGLJ0838.8–2829 source were carried out on 2016 January 14 with the 2.1-m telescope of the Observatorio Astrofísico Guillermo Haro (OAGH) in Cananea, Mexico. The telescope is equipped with the Cananea Near Infrared Camera (CANICA), which provides an image scale of 0.32 arcsec/pixel, and covers a field of view of about 4×4 arcmin². CANICA observed with the *H* filter starting at 08:03:46 UT for an exposure of 60 s, and with the *K* filter start-

Table 2. Results of the fits of the *XMM-Newton* EPIC pn spectra of RX J0838–2827 at the minimum and maximum of the ~ 54.7 ks beat modulation. For each beat phase the spectra of both observations were fitted together to the TBABS*PCFABS*(MEKAL+MEKAL+MEKAL) model, with all normalizations free to vary. Uncertainties are quoted at the 90% confidence level for a single parameter of interest.

Parameter	Minimum	Maximum
N_{H}^a (10^{20} cm $^{-2}$)	$2.3^{+0.8}_{-0.6}$	2.4 ± 0.5
Pcfabs N_{H} (10^{24} cm $^{-2}$)	$1.8^{+1.3}_{-0.7}$	$2.9^{+1.8}_{-1.1}$
Covering fraction (%)	72^{+18}_{-24}	67^{+23}_{-30}
A_{Z} ($A_{\text{Z},\odot}$)	$0.2^{+0.2}_{-0.1}$	0.37 ± 0.1
kT_1 (keV)	0.22 ± 0.03	0.12 ± 0.02
Normalization ^b	$7.1^{+1.8}_{-0.3}$ (obs 1)	$30.2^{+6.7}_{-1.8}$ (obs 1)
	$5.2^{+1.4}_{-0.3}$ (obs 2)	$21.3^{+4.7}_{-1.3}$ (obs 2)
kT_2 (keV)	1.03 ± 0.14	1.00 ± 0.08
Normalization	$8.9^{+2.4}_{-0.6}$ (obs 1)	$4.2^{+1.7}_{-0.4}$ (obs 1)
	$6.5^{+1.9}_{-0.5}$ (obs 2)	$7.6^{+0.6}_{-0.3}$ (obs 2)
kT_3 (keV)	$11.7^{+7.3}_{-2.6}$	$14.8^{+1.2}_{-0.9}$
Normalization	$38^{+7.7}_{-1.7}$ (obs 1)	161^{+31}_{-8} (obs 1)
	$34^{+6.9}_{-1.6}$ (obs 2)	130^{+13}_{-6} (obs 2)
Absorbed flux ^c	2.3 ± 0.7 (obs 1)	10.3 ± 1.0 (obs 1)
	2.0 ± 0.7 (obs 2)	8.4 ± 1.2 (obs 2)
χ^2_{ν} (dof)	1.09 (213)	1.06 (541)
Null hypothesis probability	1.5×10^{-1}	1.3×10^{-1}

^a The abundances are those of Wilms et al. (2000). The photoelectric absorption cross-sections are from Verner et al. (1996).

^b Normalizations are in units of 10^{-4} cm $^{-5}$.

^c Fluxes are in units of 10^{-12} erg cm $^{-2}$ s $^{-1}$ and in the 0.2–10 keV energy range.

ing at 08:13:56 UT for an exposure of 30 s. The counterparts to the three X-ray sources were detected in both bands, with magnitudes: $H = 17.16 \pm 0.2$ and $K = 16.18 \pm 0.09$ for RX J0838–2827, $H = 17.0 \pm 0.2$ and $K = 16.7 \pm 0.1$ for XMM J083850.4–282759, and $H = 17.0 \pm 0.2$ and $K = 16.2 \pm 0.1$ for XMM J083842.7–283832. The zero points were 20.18 and 21.34 for the H and K filters, respectively (see also Fig. 5). Neither XMM J083850.4–282759 nor XMM J083842.7–283832 have a counterpart in the Two-Micron All-Sky Survey (2MASS) catalogue⁵.

3.2 XMM-Newton Optical Monitor

The *XMM-Newton* Optical/UV Monitor Telescope (OM; Mason et al. 2001) was configured in ‘Image’ mode throughout the first observation and used the V filter, providing a wavelength coverage within the 5100–5800 Å range. A total of 10 exposures were acquired, resulting in a total dead-time-corrected on source exposure time of about 45.8 ks (the first 5 exposures had a duration of about 5 ks each, whereas all the following 5 exposures lasted about 4.16 ks). The OM acquired the data in both ‘fast-window’ and ‘Image’ modes during the second observation, using the B filter (centred on 4392 Å and with band pass between 3800 and 5000

Table 3. Log of IAC80 imaging observations and results of photometry of the source XMM J083850.4–282759. Magnitudes are in the Vega system and are not corrected for Galactic extinction.

mid exp time (MJD)	texp (s)	Filter	mag
<i>XMM J083850.4 – 282759</i>			
57367.21949	18×300 s	B	22.20 ± 0.17
57367.22329	18×300 s	R	19.83 ± 0.12
57368.20316	16×300 s	B	21.86 ± 0.13
57368.20314	16×300 s	R	19.71 ± 0.12
57392.11902	11×300 s	B	21.31 ± 0.11
57392.12283	12×300 s	R	19.55 ± 0.12
57415.07043	15×300 s	B	> 21.5 (3σ U.L.)
57415.06659	16×300 s	R	20.08 ± 0.15

Å). No useful analysis could be carried out on the data sets taken in fast-window mode, because the source fell outside the small square window (23×23 px 2) of the instrument. We then focused on the data in the ‘Image’ mode, which consist of 15 4.4-ks frames for a total exposure of 66.0 ks.

We extracted the background-subtracted photometric data using the OMICHAIN processing pipeline with the default parameter settings, as recommended by the SAS threads⁶. Only one source is detected within the error circle of the X-ray position of RX J0838–2827, for which we show the light curve in Fig. 2. The optical emission clearly shows a variability correlated with the X-ray emission, around an average value of 17.77 mag in the V filter and 17.53 mag in the B filter (expressed in the Vega photometric system). XMM J083850.4–282759 and XMM J083842.7–283832 are clearly detected in the second, longest observation. Interestingly, the former exhibits a flare ($B \sim 19.5$ mag) simultaneous to that observed in the soft X-ray flux (see Section 6 and Fig. 6). No significant flaring episodes are instead observed in the data sets of the latter source.

3.3 IAC 80

Optical imaging observations of the field of 3FGL J0838.8–2829 were carried out with the IAC-80 telescope on 2015 December 11 and 12 (MJD 57367 and 57368), on 2016 January 5 and 28 (MJD 57392 and 57415). A set of exposures lasting 300 s were obtained on each night alternating the B and R filters. Image reduction was carried out by following the standard procedures: subtraction of an averaged bias frame, division by a normalized flat frame. Astrometry was performed using the USNOB1.0⁷ catalogue. Aperture photometry was made with the PHOTOM software part of the STARLINK⁸ package. The photometric calibration was done against the APASS catalogue. In order to minimise any systematic effect, we performed differential photometry with respect to a selection of local isolated and non-saturated standard stars. In Fig. 5 we show two images taken on December 11 in the B and R bands. We detected the three X-ray sources in both bands. RX J0838–2827 displays significant variability matching the periodicity observed in the X-rays and with the OM. The mean

⁶ See http://xmm2.esac.esa.int/sas/8.0.0/documentation/threads/omi_thread.html

⁷ <http://www.nofs.navy.mil/data/fchpix/>

⁸ <http://starlink.eao.hawaii.edu/starlink>

⁵ See <http://www.ipac.caltech.edu/2mass/releases/second/>.

magnitudes are $B \sim 17.7$ and $R \sim 17.0$. The other two X-ray sources are detected with mean magnitudes $B \sim 21.7$ and $R \sim 19.8$ for XMMJ083850.4–282759 and $B \sim 20.0$ and $R \sim 19.01$ for XMMJ083842.7–283832. This last source shows a steady optical flux. On the other hand, some variability is revealed in XMMJ083850.4–282759 both in the single R -band exposures (in the B -band the source is too faint to be detected in the single frames) and among the different nights (Tab. 3).

3.4 ESO NTT observations

We performed optical spectral observations of RXJ0838–2827 on 2016 February 27 with the ESO-NTT telescope (ID: 296.D-5034; PI: Rea). The observations were carried out with the EFOSC2 spectrograph using grism 19 and a slit aperture of $1''$, covering the spectral range 4445–5110 Å with resolution $R \sim 2500$ (0.67 Å pixel $^{-1}$). A total of 29 spectra, each one lasting 550 s, were obtained in the time interval from 01:05:30 UT–06:55:33 UT. Data were reduced using standard procedures using the ESO-MIDAS package for bias subtraction and flat-field correction. Wavelength calibration was carried out using helium-argon lamps. We show in Fig. 7 and 8 the acquired spectra.

4 RADIO OBSERVATIONS

We observed the field around 3FGL J0838.8–2829’s position using the *Australia Telescope Compact Array* in the 1.5A configuration at 5.5 and 9 GHz with a bandwidth of 2 GHz at each frequency (TOO project CX339). The observations were carried out on 2015 December 4 from 12:46 to 20:28 (UT), with a total integration time of 6.2 hours. We used PKS B1934–638 as the flux calibrator and 0858–279 as the phase calibrator. After standard calibration and imaging with MIRIAD (Sault, Teuben & Wright 1995), we ran one round of phase self-calibration and one round of complex amplitude self-calibration. The resulting images (with robust=0.5 weighting) have an rms of 6.3 and 7.1 μ Jy at 5.5 and 9 GHz (see Fig. 5). The synthesised beam sizes are 11.1 arcsec by 2.3 arcsec (P.A. -6.5°) and 7.35 arcsec by 1.75 arcsec (P.A. -6.5°), respectively. We detect no radio source at the position of RXJ0838–2827 or at those of the other two X-ray sources down to a 3σ upper limit of ~ 20 μ Jy at both frequencies. Other unidentified sources are detected in the field (see Fig. 5 and 6).

5 RX J0838–2827: DATA ANALYSIS AND RESULTS

5.1 X-ray timing analysis

The light curves of the two *XMM-Newton* observations of RXJ0838–2827 reveal substantial variations with time (see Fig. 2 9 and 10). A strong modulation with period of ≈ 1.6 h (6 ks) is easily recognisable, with another ample modulation on a time scale approximately ten times longer. Since the longer modulation is present in both the *XMM-Newton* observations, it might be recurrent as well, with a characteristic time ≈ 15 h (54 ks).

The modulation pattern is rather complex and we lack a prior knowledge of its origin, so we adopted the following approach to model it: we first fit to the data of each observation independently two sinusoidal function. Since the fit was not satisfactory, we tested the addition of further sinusoidal components with periods free to assume any value (i.e., we did not forced them to be integer factors

or multiples of P_1 or P_2). In other words, we modelled the data the data with a Fourier sum of the form

$$A_0 + \sum_{k \geq 2} A_k \sin\left(\frac{2\pi(x - \phi_k)}{P_k}\right). \quad (1)$$

Of course, as k increases, the adherence of the model to the data improves, but we found that the simplest function that fits reasonably well the data, adequately describing the occurrence of minima and maxima, and the overall shape of the curve, requires three sin components ($k = 3$). By fitting this purely phenomenological model to the data, we derived the values $P_{1, \text{XMM1}} = 1.644 \pm 0.005$ h (5.92 ± 0.02 ks), $P_{2, \text{XMM1}} = 15.0 \pm 0.5$ h (54.2 ± 2.0 ks), and $P_{3, \text{XMM1}} = 1.491 \pm 0.008$ h (5.37 ± 0.03 ks) for the first *XMM-Newton* observation, and $P_{1, \text{XMM2}} = 1.650 \pm 0.003$ h (5.94 ± 0.01 ks), $P_{2, \text{XMM2}} = 15.2 \pm 0.2$ h (54.7 ± 0.7 ks), and $P_{3, \text{XMM2}} = 1.477 \pm 0.003$ h (5.34 ± 0.01 ks) for the second. The measures of P_1 and P_2 are not precise enough to bootstrap a unique coherent solution that links the two *XMM-Newton* data sets (in particular, there is a strong alias degeneracy in P_1).

To better assess the situation, we inspected a Fourier transform of the second *XMM-Newton* observation computed using the pn data which allow the highest timing resolution (bin time: 47.66 ms). Several peaks can be observed in the power spectral density (PSD) plot (Fig. 9). Two strong close but distinct peaks (their separation is larger than the intrinsic Fourier resolution, $\Delta\nu \simeq 1.4 \times 10^{-5}$ Hz) are located at $\Omega = 1.69(5) \times 10^{-4}$ Hz (the central frequency was measured by the fit with a Gaussian, and is equivalent to 1.64 ± 0.04 h) and $\omega = 1.88(9) \times 10^{-4}$ Hz (1.47 ± 0.08 h), and they clearly correspond to $P_{1, \text{XMM2}}$ and $P_{3, \text{XMM2}}$, respectively. An even higher peak is present at $\sim 1.87 \times 10^{-5}$ Hz (14.88 h) and is consistent with $P_{2, \text{XMM2}}$. The fact that the central frequency of this peak differs by less than 2% from the quantity $\omega - \Omega$, strongly suggests that it is the sideband (beat) frequency between Ω and ω , which we then interpret as the orbital and spin period of RXJ0838–2827, respectively (see Sect. 7.1). Many other peaks that stand out above the significance threshold in the PSD can be interpreted as combinations of Ω and ω . The most prominent ones are labelled in Fig. 9. The PSD computed from the first observation is similar (Fig. 9), but only one peak is present around the frequencies of Ω and ω . This is to be expected in any case, because of the lower intrinsic resolution of this observation ($\sim 2.2 \times 10^{-5}$ Hz), but the slightly asymmetric shape of the peak, which is unbalanced towards Ω , suggests that it may actually be a blend of two peaks.

5.2 X-ray spectral analysis

We started the spectral analysis by fitting together the pn averaged-spectra of the two observations to a set of different models: power laws, bremsstrahlung, blackbodies, an accretion disc consisting of multiple blackbody components (DISKPB), a model describing the emission produced by a thermal distribution of electrons which Compton up-scatter soft seed X-ray photons (NTHCOMP; Zdziarski, Johnson & Magdziarz 1996; Zycki, Done & Smith 1999), and a model reproducing the emission by hot diffuse gas (MEKAL). We accounted for photoelectric absorption by the interstellar medium along the line of sight through the TBABS model, and adopted the photoionisation cross-sections from Verner et al. (1996) and the chemical abundances from Wilms, Allen & McCray (2000). None of the above-mentioned models gave a satisfactory description of the data.

Much more acceptable results were achieved when fitting the

data to models consisting of several MEKAL components at different temperatures.

Given the strong flux variability in both observations, as well as the variable hardness ratio we observed comparing light-curves in different energy bands (see Fig. 2), we performed the spectral analysis comparing time-resolved spectra, which we believe yielding more meaningful results for this variable source (see Fig. 11). We first analysed the spectra during the peak of the 54.7 ks beat modulation for both observations, and then we further extracted spectra for the maximum and minimum of the shorter 5.9 ks periodicity. Fitting with a series of 3MEKAL models we find good results for all the spectra we considered. Despite cycle-to-cycle variations observed in the hardness ratio in both observations (see Fig. 2), we did not find significant variability (within the 90% parameter uncertainties) when fitting the 1.64 h phase-dependant spectra in any of the two observations nor merging them in phase (most probably due to the low number of counts we could collect at the minimum phases). However, we find significant changes in the spectra collected at the maximum and minimum of the long-term beat period of 54.7 ks. Best-fitting parameters for the absorbed 3 MEKAL model are listed in Table 2, whereas the spectrum and the best-fitting model are shown in Fig. 11. It is evident how the relative significance of the three different hot plasmas changes along the beat modulation, with the coolest component at an effective temperature of ~ 0.2 keV being more prominent at the maxima of the beat ($\omega - \Omega$) modulation. This might be due to accretion switching to the other (lower) pole-switching at half beat cycle, and the other pole possibly not totally visible (depending on the colatitude and the inclination angles).

The source absorbed $0.2\text{--}10$ keV flux during the first observation was 7.6×10^{-12} erg cm $^{-2}$ s $^{-1}$, a factor of ~ 1.2 larger than measured during the second observation, about 1.5 months later. The derived absorption column density, $N_{\text{H}} \sim 1.6 \times 10^{20}$ cm $^{-2}$, is about one order of magnitude lower than the estimated total Galactic value in the direction of the source ($\sim 2 \times 10^{21}$ cm $^{-2}$; Willingale et al. 2013), implying a close-by location of this source within our Galaxy.

5.3 High resolution X-ray spectra

We used the first order RGS spectra of both observations to search for narrow features. Two prominent narrow emission lines are clearly detected in both observations, plus a few weaker lines were also significantly detected. All the observed lines were well fitted by the multiple hot ionised plasma models (MEKAL) used to fit the EPIC-pn spectra (see Sec 5.2). In particular, we identified in the RGS spectrum of the second XMM observations the lines corresponding to: Mg XII H-like Lyman α at 8.42 Å, Ne X H-like Lyman α at 12.13 Å, Ne IX He-like w-resonance at 13.45 Å, Oxygen VIII H-like Lyman α (18.97 Å) and β (16.01 Å), and triplet of the He-like Oxygen VII centred at 21.60 Å (see Fig. 3).

To test possible signs of variability among the most prominent lines between the maximum and minimum of the 1.64 h X-ray modulation, we fitted the two phase-resolved RGS data (merging data of both observations) with an absorbed power-law plus two Gaussian functions with energies fixed at the 21.60 Å (0.57 keV) and 18.97 Å (0.65 keV), and keeping all spectral parameters forced to be the same except for the normalizations. The equivalent width of the Oxygen VII line is 64^{+34}_{-17} eV in the RGS spectrum at the maximum of the modulation, while only an upper limit of < 62 eV (at 90%) could be derived in the RGS spectrum at the minimum of the 1.64 h modulation. Similarly, the equivalent width of the Oxygen

VIII line is 33^{+6}_{-13} eV in the RGS spectrum at the maximum, and we derived an upper limit of < 69 eV at the minimum. No significant phase-dependent variability is detected in the emission lines, and the Oxygen lines are well reproduced by the lowest temperature (~ 0.2 keV) MEKAL component.

5.4 Long-term X-ray light curve

The long-term light curve of RX J0838–2827 presented in Fig. 4 was extracted after fitting all *Swift*/XRT spectra with an absorbed power law model with the absorption column density fixed to the value derived from the fits of the *XMM-Newton* data (the lower counting statistics of the XRT data precluded a more detailed modeling as done for the EPIC data sets). We obtained $\chi^2_{\nu} = 1.13$ for 203 dof. We deem that the large scatter in the flux values clearly visible in the figure (especially in the most recent observations) are likely due to the fact that the short exposures of the observations (see Table 1) have sampled the source ≈ 54 -ks cycle at random phases. However, the long-term light-curve indicates that no substantial variation is observed in the source X-ray flux across the past 6 years.

5.5 Results from optical observations

The optical spectra of RX J0838–2827 (see Fig. 7) are dominated by prominent emission lines such as H β , He I (λ 4471, 4919, 5015), He II (λ 4686) and possibly bowen-blend N III (λ 4634-41)/C III (λ 4647-50). No significant absorption feature is detected, in agreement with the findings of Masetti et al. (2013).

Among the emission features detected, H β is the most prominent. In the individual spectra we found no evidence for double-horned profile (usually interpreted as a signature for an accretion disc). The line profile is however clearly asymmetric. We tried to model the H β line with a set of multiple Gaussians and found a reasonable solution using three Gaussians plus a constant: in Fig. 12 we show the velocities of the these three components. The H β line consists of a narrow (FWHM = 11 Å ~ 680 km s $^{-1}$) a broad (FWHM ~ 30 Å ~ 1900 km s $^{-1}$) and a very narrow (FWHM ~ 6 Å ~ 370 km s $^{-1}$) components, the latter statistically detected only in a few spectra (see Fig. 12). The large width of the broad component indicates it probably forms in the magnetically confined accretion flow, the narrow one likely forming in the ballistic trajectory of the accretion stream and the very narrow component from the irradiated hemisphere of the donor star. These component are generally observed in Polars and particularly in asynchronous polars (see eg. Schwarz et al. 2005).

Fits to the radial velocity measurements for each lines have been first attempted by using a constant and a sinusoidal function (see Fig. 12). We here use the radial velocities as derived for the H β narrow and broad components, and the radial velocity measures of the other emission lines obtained through single-gaussian. A fit with a constant plus a single sinusoidal function never provided a satisfactory result in terms of reduced χ^2 . We then tried a different approach and fit each radial velocity curve with the function of eq.(1) with $k = 3$, $P_2 = P_{2, \text{XMM2}} = 15.2$ h and $P_3 = P_{3, \text{XMM2}} = 1.47$ h (fixed); see Fig. 12). We measured the following P_1 periods: 1.75 ± 0.05 h (H β , first velocity, $\chi^2_{\nu}/\text{dof} = 1.40/21$), 1.72 ± 0.05 h (H β , second component $\chi^2_{\nu}/\text{dof} = 0.71/21$), 1.53 ± 0.01 h (He I 4480 Å, $\chi^2_{\nu}/\text{dof} = 0.67/20$), 1.48 ± 0.01 ks (He I 4920 Å, $\chi^2_{\nu}/\text{dof} = 4.53/20$), and 1.66 ± 0.04 h (He I 5020 Å, $\chi^2_{\nu}/\text{dof} = 0.61/20$). From the first H β

component, we were able to derive some constraints also on the amplitudes: $A_0 = -(4 \pm 1) \times 10^2 \text{ km s}^{-1}$, $A_1 = (4.2 \pm 0.5) \times 10^2 \text{ km s}^{-1}$, $A_2 = (6.3 \pm 1.6) \times 10^2 \text{ km s}^{-1}$, $A_3 = (1.7 \pm 0.3) \times 10^2 \text{ km s}^{-1}$. This multi-sinusoidal representation indicates that besides the orbital modulation (with amplitude of about 400 km/s - see Fig. 8) a fast moving high velocity component shifting from a maximum of about +1100 to -1200 km/s may be due to the ballistic accretion stream which changes also along the beat cycle.

For each spectrum, we also measured the equivalent width (EW) of the H β and He I emission lines. We found evidence for significant EW variability, as observed also for the optical flux measured by our IAC-80 and XMM-Newton/OM observations (see e.g. Fig. 2 right panel-bottom). The EW of all emission features seems to be modulated over the 1.64 h and 15.2 h periodicities, the orbital period and the longterm beat periodicity (see Fig. 12). Since the spectral lines EW is a proxy of the source optical flux, such findings are not unexpected.

The NTT data consist of short exposures of ~ 500 s each, spanning only a few cycles of the orbital modulation at 1.64 h, for a total observation time of ~ 6 h (see Fig. 8 and 12). For this reason, when fitting them with the multicomponent models used in Section 5.1, we held P_2 and P_3 at the XMM-Newton values measured during the second observation.

Because of the very sparse sampling, long data gaps, and short exposures, the IAC80 photometric data are of limited help in constraining the model and/or refining the periods. So, we fixed $P_1 = P_{1, \text{XMM2}} = 1.64$ h and $P_3 = P_{3, \text{XMM2}} = 1.47$ h (the 3-sin model provides also for the IAC80 data a much better fit than the 2-sin model). In this way, we found for the long modulation $P_{2, \text{B}} \simeq 15.1$ h from the B-filter data and $P_{2, \text{R}} \simeq 15.2$ h from the R-filter data.

6 THE OTHER TWO X-RAY SOURCES IN THE FIELD: XMM J083850.4–282759 AND XMM J083842.7–283832

Of the two additional X-ray sources within the field of the mCV, XMM J083850.4–282759 can be identified in the USNO A2 Catalogue at RA = $08^{\text{h}}38^{\text{m}}50^{\text{s}}.4$ Dec = $-28^{\circ}27'56''.7$ (J2000.0) with $R = 18.0$ and $B = 20.6$, and XMM J083842.7–283832 is catalogued in the Wide-field Infrared Survey Explorer (WISE⁹; Wright et al. 2010) as WISE J083842.77–282830.9. Its position is RA = $08^{\text{h}}38^{\text{m}}42^{\text{s}}.777$, Dec = $-28^{\circ}28'30''.99$ (J2000.0), and it has optical magnitudes in the USNO B1 Catalog of $B1 \sim 19.66$, $B2 \sim 18.49$, $R2 \sim 19.02$ and $I \sim 18.01$.

We carried out a detailed X-ray spectral and timing analysis for both sources. XMM J083850.4–282759 appears to be rather variable in the X-rays (see the values for the root mean square fractional variation and fluxes in Table 4), and showed a ≈ 600 -s long flare with a structured morphology in the second observation (see Fig. 6).¹⁰ For both sources, spectra of the two observations were fitted together with an absorbed power law model, and the absorption column density was tied up across the data sets (for the case of XMM J083850.4–282759 both the flare and quiescent spectra were extracted). Best-fitting parameters are reported in Table 4. The steady emission of XMM J083850.4–282759 is brighter during the second observation by a factor of ~ 5 , while showing flaring

episode, and its spectrum slightly harder than in the first observations. On the other hand, XMM J083842.7–283832 appears to be steady during the X-ray observations.

7 DISCUSSION

Our multi-band study of the field of 3FGL J0838.8–2829 revealed two variable X-ray sources: RX J0838–2827 and XMM J083850.4–282759. We discuss here our results and the nature of these two intriguing objects.

7.1 RX J0838–2827: an asynchronous magnetic CV

RX J0838–2827 has shown to be highly variable on short and long timescales with a complex behaviour. In particular, we detect two dominant periodicities at 1.64 h and 15.2 h in the power spectrum as well as a 1.47 h peak which is about half the power of the 1.64 h one (see Fig. 9). These are remarkably correlated as the long periodicity is found precisely at the beat between the two shorter ones. These characteristics and their timescales are uncommon to other variable X-ray sources, as the LMXBs, but are consistent with those observed in CVs of the magnetic type. The close 1.64 h and 1.47 h periods if interpreted as the binary orbital period and the rotational period of the WD give $P_{\text{spin}}/P_{\text{orb}} \sim 0.90$. The location of RX J0838–2827 in the spin-orbit period plane of mCVs is shown in Fig. 13, and strongly suggests this system is close to, but not right at, synchronism. We therefore have identified either one of the slowest rotating IP (see also Coti Zelati et al. 2016) or a highly de-synchronised polar. Other weakly asynchronous magnetic systems with similar spin-orbit ratios are Paloma, $P_{\text{spin}}/P_{\text{orb}} \sim 0.93$ (Schwarz et al. 2007), and IGR J19552+0044 with $P_{\text{spin}}/P_{\text{orb}} \sim 0.82$ (Bernardini et al. 2013). The other four confirmed asynchronous polars are only very slightly de-synchronised ($P_{\text{spin}}/P_{\text{orb}} > 0.98$; see also Fig. 5).

Folding the X-ray light curve at the 1.64 h period (see Fig. 10; or at the 1.47 h too) the shape is almost flat-topped closely resembling those observed in the strongly magnetic polar systems (Cropper 1990), where the modulation is due to self-occultation of the accretion pole as the WD rotates. The spectrum is thermal, requiring three optically thin emitting regions with temperatures 0.2 keV, 1 keV and 14 keV highly absorbed by dense partial ($\sim 70\%$) covering material ($\sim 2 \times 10^{24} \text{ cm}^{-2}$). The spectral characteristics also closely resemble those observed in magnetic CVs, both polars and IPs (e.g. Ramsay & Cropper 2004; Bernardini et al. 2012). The lack of spectral changes (except for the normalization) at the ~ 1.6 – 1.5 h periods is consistent with the spectral behaviour seen in polars due to the self-occultation of the main accreting pole, while IPs instead show highly variable absorption along the spin period. Furthermore, and differently from many polars, RX J0838–2827 does not show any evidence of a soft optically thick (blackbody) component. This component originates from reprocessing of hard X-rays and cyclotron radiation at the WD polar regions. The presence of this component has been for a long time considered ubiquitous in polars, until new observations have shown that this is not necessarily so (Ramsay & Cropper 2004). On the other hand, IPs have typically hard spectra, and the recent identification of a soft blackbody component in about 30% of these systems (Anzolin et al. 2008; Bernardini et al. 2012) does not make these characteristics an identifying feature of one or the other type of mCVs.

The long observations acquired with XMM-Newton have allowed an unprecedented monitoring of RX J0838–2827 revealing

⁹ See <http://wise2.ipac.caltech.edu/docs/release/allsky/>.

¹⁰ A careful analysis of the background light curve excludes a background flare origin

Table 4. Positions, pn count rates and spectral fit results for the X-ray sources within the error box of 3FGL J0838.8–2829. The source numbering is as shown in Fig. 6. The pn spectra were fitted to the TBABS*PEGPWLW model in the 0.2–10 keV energy range. Uncertainties are quoted at a 90% confidence level for a single parameter of interest, whereas upper limits are given at the 3σ confidence level. (q) and (f) refer to the quiescent and flaring states of XMM J083850.4–282759, respectively, during the second observation.

Source	Obs ID	RA (J2000.0)	Dec	Count rate ^a ($\times 10^{-3} \text{ s}^{-1}$)	rms ^b	N_{H} (10^{21} cm^{-2})	Γ	Unabs flux ^a ($10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$)	χ^2_{ν} (dof)
XMM J083850.4–282759	0764420101	08 ^h 38 ^m 50 ^s .40	–28° 27′ 59″.04	16.7 ± 0.9	0.37 ± 0.06	0.3 ± 0.2	1.6 ± 0.2	8 ± 1	0.97 (63)
	0790180101			74 ± 2 (q)	1.23 ± 0.07	tied	1.15 ± 0.08 (q)	38 ± 3 (q)	
				1223 ± 82 (f)	tied	1.3 ± 0.2 (f)	550 ± 90 (f)		
XMM J083842.7–283832	0764420101	08 ^h 38 ^m 42 ^s .72	–28° 28′ 32″.52	22.1 ± 0.9	< 0.18	1.1 ± 0.3	1.9 ± 0.2	11.1 ± 0.9	0.94 (27)
	0790180101			14.8 ± 0.8	< 0.17	tied	2.1 ± 0.2	8.6 ± 0.9	

^a In the 0.2–10 keV energy range.

^b Root mean square fractional variation, determined by applying the LCSTATS (v. 1.0) tool of XRONOS on the 0.2–10 keV background-subtracted and exposure-corrected light curves from the pn camera with a binning time of 400 s.

a remarkable long-term 15.2 h X-ray modulation which repeats itself in the two observations taken 1.5 months apart that we naturally interpret as the beat period between the 1.64 h and 1.47 h periodicities. The spectrum changes mainly in the normalizations of the optically thin components and in the temperature of the warmer plasma but not in the partial absorption components. This further strengthens our interpretation that this variability is not the orbital period since the orbital X-ray variability in the asynchronously rotating IPs is found to be due to absorbing material fixed in the binary frame such as the hot-spot at the rim of an accretion disc or eclipses in high inclination systems (see Parker et al. 2005). The optical radial velocities of emission lines and the optical light also give evidence for the presence of such long variability implying that it does not arise from X-ray irradiation on an accretion disc as seen in many IPs (see Warner 2003). The presence of this modulation in both X-rays and optical implies a diskless configuration (Wynn & King 1992), in which the accreting material alternates onto the WD poles. In non-synchronous systems, the flow rotates around the magnetic field of the WD on a time-scale of the beat period. This has the effect that the flow will be preferentially directed onto the first (main) pole and then to the other, each half beat cycle. The weaker modulations at the 1.47 h and 1.64 h periods at about half the 15.2 h beat period indicates that there is pole switching in RX J0838–2827. In particular diskless models predict the dominance of the beat and of the orbital frequencies and a much lower or even undetectable spin peak in the power spectrum in configurations of moderately high values of the binary inclination angle (i) and magnetic colatitude (m) provided that $i + m > 90^\circ + \beta$, where β is the pole-cap opening angle (Wynn & King 1992). The presence of lower peaks at sidebands depends on pole-switching effect that modulates the light curve. Furthermore the power of the spin peak is dependent on the degree of asymmetry, with strong asymmetries producing weaker spin pulsations (Wynn & King 1992). This appears to be the case in RX J0838–2827. Additionally the slope of the beat modulation depends on the transition times of accretion flipping onto the two poles and thus the almost smooth variability at the 15.2 h period would indeed indicate that the stream from the donor star does not switches between the two poles instantaneously. We furthermore note that a large β tends to produce either sinusoidal or flat-topped light curves indicating that accretion in RX J0838–2827 occurs over wide areas of the WD poles. Furthermore, the faint phase of the modulation at the 1.64 h (or 1.47 h) period (see Fig. 10) observed at beat maximum lasts ~ 0.3 . This can be used to estimate the ranges of binary inclination i and magnetic colatitude m (Cropper 1990) as: $\cot i = \cos(\pi \delta\phi) \times \tan(m)$, where $\delta\phi$ is the length of the faint phase. Also, the lack of X-ray eclipses restricts $i \lesssim 75^\circ$. We then find $24^\circ \lesssim m \lesssim 59^\circ$ for $10^\circ \lesssim i \lesssim 75^\circ$. A detailed modelling of accretion region is however beyond the

scope of this work and will be subject of a future publication. The flipping accretion configuration are likely occurring also in other asynchronous systems such as Paloma (Schwarz et al. 2007), BY Cam (Piirola et al. 1994) and CD Ind (Ramsay et al. 1999), and in V1432 Aql (Mukai et al. 2003). In the latter system, the hard X-ray emission is found to be dominated by the orbital variability along the beat cycle whilst the X-ray spin variability is almost undetected in a long *BeppoSAX* observation (Mukai et al. 2003) and with complex structure in *XMM-Newton* and *RXTE* observations (Rana et al. 2005).

The short orbital period of RX J0838–2827 puts this mCV below the 2–3 h orbital period gap of CVs, where most polars are found (Ferrario, de Martino & Gaensicke 2015; see also Fig. 13). An orbital period of 1.64 h (98 min) is close to that of the IP EX Hya which is also (but not so weakly) de-synchronised ($P_{\text{spin}} \sim 67$ min). Its spin-orbit period ratio ($P_{\text{spin}}/P_{\text{orb}} = 0.68$), similar to the few IPs found below the gap, could indicate that these systems will never reach synchronism (see Norton et al. 2004). Looking at the other two weakly de-synchronised mCVs, Paloma and IGR J19552+0044, they are both short period systems, the former with a 157 min orbital period falls right inside the orbital period gap and the second with a 101 min orbital period is below the gap. They are likely in a transition stage before reaching synchronism and thus believed to be polar progenitor candidate. On the other hand, the four asynchronous polars are instead less de-synchronised. Synchronisation timescales of ~ 170 yr and ~ 200 yr were found in V1500 Cyg (Harrison & Campbell 2016) and V1432 Aql (Staubert et al. 2003), respectively, suggesting that the cause was likely a nova explosion in the past. The lack of nova shells around the 4 systems is not however a clue against the nova interpretation, since many old novae do not show such relics either (Pagnotta & Zurek 2016).

Therefore, given the closer similarities of spin-to-orbit period ratios with Paloma and IGR J19552+0044, rather than to the four asynchronous polars, we suggest RX J0838–2827 is in a similar evolutionary stage. A spectro-polarimetric study will prove tight constraints on the magnetic field and geometry of this new interesting mCV and its evolution.

7.2 XMM J083850.4–282759: a candidate transitional millisecond pulsar?

The identification of RX J0838–2827 as an asynchronous polar argues against its identification as a possible counterpart of the *Fermi*-LAT gamma-ray source 3FGL J0838.8–2829, as emission at \sim GeV energies has never been reported from CVs, so far. On possibility might be that the gamma-ray emission is associated to

Table 5. Main properties of the slightly desynchronised mCVs currently known (in order of increasing degree of asynchronism).

System	P_{spin} (s)	P_{orb} (s)	$P_{\text{spin}}/P_{\text{orb}}$	P_{beat} (hr)
V 1432 Aquilae	12 150	12 116	1.002	1 213.4
CD Ind	6 579	6 649	0.989	175.2
BY Camelopardalis	11 961	12 089	0.989	348.5
RX J0838–2827	5 340	5 940	0.898	15.2
Paloma	7 800	9 360	0.833	13.0
IGR J19552+0044	4 960	6 100	0.813	7.3
EX Hydrae	4 022	5 895	0.682	3.5

* The listed values are taken from: V 1432 Aquilae: Mukai et al. (2003); CD Ind: Myers et al. (2017); BY Camelopardalis: Silber et al. (1997); RX J0838–2827 (this study); Paloma: Joshi et al. (2016); IGR J19552+0044: Bernardini et al. (2013); EX Hydrae: Mauche et al. (2009).

one of the two faint X-ray sources detected in the *Fermi*-LAT error circle.

As reported in Tab. 4, the absorption column density derived from the X-ray spectrum of XMMJ083850.4–282759 is $N_{\text{H}} \sim 3 \times 10^{20} \text{ cm}^{-2}$, about one order of magnitude lower than the total Galactic value in the direction of the source. Assuming that the full observed column density contributes to the optical/NIR absorption, we estimate a colour excess $E(B - V) = 0.05 \text{ mag}$ (using the $N_{\text{H}}/E(B - V)$ conversion of Predehl & Schmitt 1995). From the results of our optical and NIR photometry (Sect. 3.1 and 3.3) we obtain for this source a de-reddened $B - R$ colour in the range 1.5–2.5 mag (taking into account the source variability and the uncertainties in the photometry) and $H - K = 0.3 \pm 0.2 \text{ mag}$, broadly consistent with a late-type K2–M5 main sequence star. Under this hypothesis, using the distance modulus, the constraints on the source distance are very broad, ranging from 0.7 (M5) to 6.2 kpc (K2).

In light of the observed colours, a coronal flare from an M-type star could be an explanation for the brightening observed in the *XMM-Newton* X-ray and optical light curves shown in Fig. 6. However, considering the inferred distance, the average X-ray flux ($\sim 8.5 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the 0.2–10 keV, excluding the flare) would translate into a luminosity $5 \times 10^{30} < L_X < 4 \times 10^{32} \text{ erg s}^{-1}$, which is too high for an active M-type star in “quiescence” (see, e.g., Stelzer et al. 2013).

An interesting possibility is the association of 3FGLJ0838.8–2829 with the X-ray variable source XMMJ083850.4–282759, as a possible transitional millisecond pulsar (Archibald et al. 2009; Papitto et al. 2013). Such sources are often found in a sub-luminous disc state characterized by enhanced X-ray variability between a high and a low flux level that differ by a factor of ~ 10 , and often show occasional flaring activity on several minutes timescales (de Martino et al. 2013; Bogdanov et al. 2015), as we observe from XMMJ083850.4–282759. A relatively bright emission is also observed at gamma-ray (de Martino et al. 2010, Stappers et al. 2014; Li et al. 2014), optical (de Martino et al. 2013; Bogdanov et al. 2015) and radio (Hill et al. 2011; Deller et al. 2015) wavelengths. Even though the absence of an ATCA radio detection from XMMJ083850.4–282759 does not favour its identification as a transitional millisecond pulsar, the non detection of a radio source might be compatible with a larger distance of the putative pulsar. In this respect, we note that the gamma-ray flux of 3FGLJ0838.8–2829 ($\sim 10^{-9} \text{ ph cm}^{-2} \text{ s}^{-1}$ in the 1–100 GeV band; Acero et al. 2015), and its average

X-ray flux ($\sim 8.5 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}$ in the 0.2–10 keV band) are ~ 10 and 100 times smaller, respectively, than those observed from the prototypical transitional millisecond pulsar PSR J1023+0038, which is located at 1.4 kpc (Deller et al. 2012), hence a larger distance would be also in line with other multi-band source characteristics. Furthermore, the low spectral index and low level of variability of the GeV emission are also inline with what observed for other binary millisecond pulsars observed by *Fermi*-LAT (Acero et al. 2015).

The detection of optical/NIR emission from a late-type companion star, as inferred from the observed colors, would suggest for a redback system observed in the radio pulsar state. The constraints on the X-ray luminosity reported above would be in agreement with such a scenario. However, without knowing the orbital period, the amount of possible irradiation effects, the companion spectral type and the system distance, all the above considerations should be considered with some caution and alternative scenarios (e.g. a black-widow system) could not be ruled out.

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8 APPENDIX

During the referee process of this paper, another work using our *XMM–Newton* data were published by Halpern, Bogdanov & Thorstensen (2017; *ApJ*, 838, 124). In this Appendix we point out a few differences between the two publications.

1) In our optical monitoring we could not detect any hint for periodicity in the optical monitoring of the transitional MSP candidate XMMJ083850.4–282759, nor the broad minimum reported by Halpern et al. (2017), but we caution, that the longest time-series we collected span

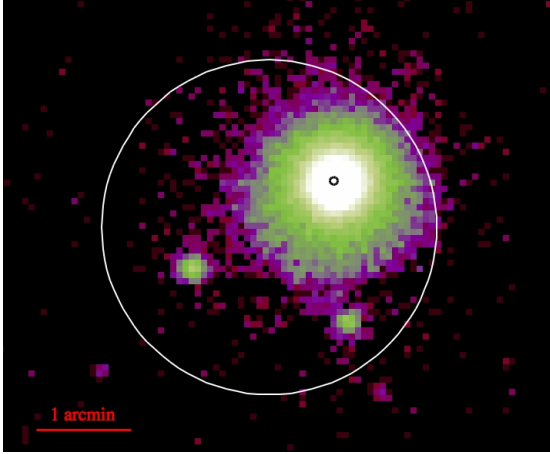


Figure 1. Field around RX J0838–2827 in the 0.2–10 keV band as observed by *XMM–Newton*. The image shows the merging of all observations and EPIC instruments, for a total exposure time of ~ 125 ks. The white circle shows the position of the Fermi-LAT source 3FGL J0838.8–2829 with a 0.03° error radius (at 68% confidence level) as reported in the 3FGL catalog. The black circle is centred on the best optical position, with an error circle of 2 arcsec (enlarging by a factor of 10 the optical positional accuracy for imaging purposes). North is up, and east is left.

over about 2.9 hours (about one half with respect to the length of Halpern et al. optical dataset).

2) By using our *XMM–Newton* data, they report different modelling/interpretation of the X-ray light curve. They assumed that a phase jump of $\sim 120^\circ$ in the ~ 1.6 h modulation happens around the flux minimum and, using only the first part of the observations, derived a spin period of ≈ 1.58 h, which is barely compatible with our value. As for the orbital period, they measured it from radial velocities of $H\alpha$ in good quality optical spectroscopic data and found 1.64 ± 0.06 h, in good agreement with our values. These differences reflect in a different interpretation on the nature of the mCV RX J0838–2827. To see whether one description of the light curve is to be preferred to the other, we analysed separately the *XMM–Newton* segments of data prior and after their suggested phase jump in both observations. The comparison was inconclusive, essentially because the orbital and spin periods derived in the two scenarios are almost identical and because not all subsets of data have enough frequency resolution to separate the Ω and ω peaks in the PSDs (see Fig. 14 for an example).

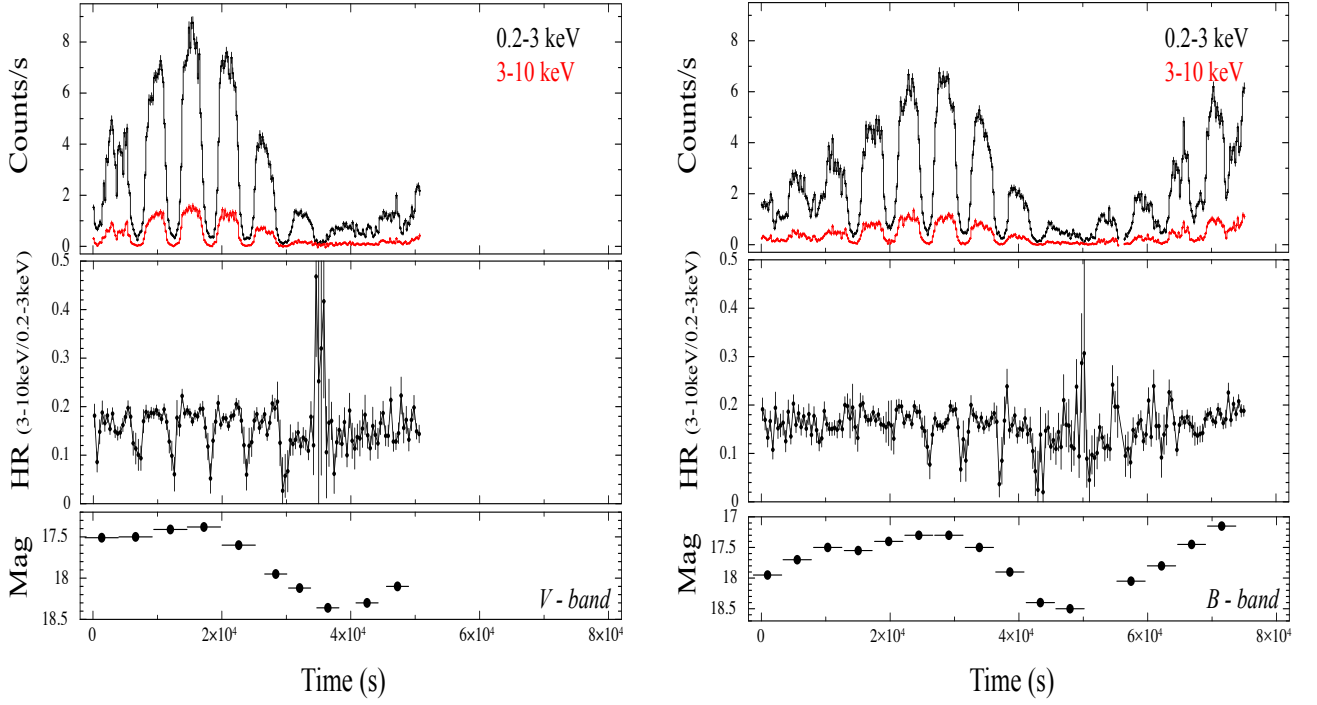


Figure 2. *XMM–Newton* EPIC (pn+MOS) lightcurves (with a binning time of 200 s) of RX J0838–2827 in two energy bands (and hardness ratios) for the first (left) and second (right) *XMM–Newton* observation. In the bottom panels we report on the *XMM–Newton* Optical Monitor magnitudes during the two observations (see text for details). Errors are smaller than the markers (~ 0.02 – 0.04 mag).

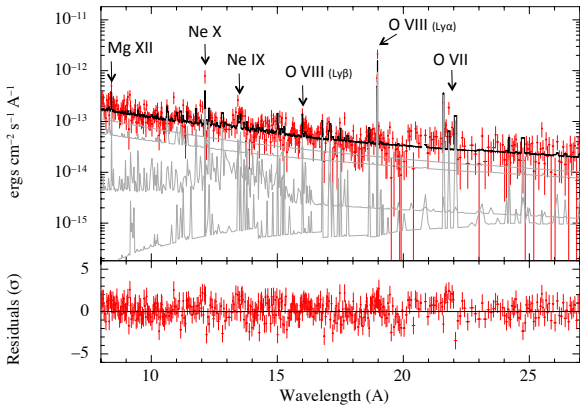


Figure 3. *XMM–Newton* RGS first order spectra of RX J0838–2827 for the second observation. Grey lines represent the three ionized plasma components of the model (black line) that best fits the pn (and RGS) data.

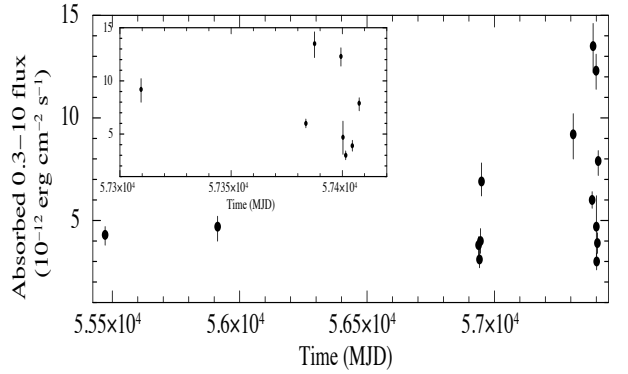


Figure 4. *Swift* XRT long-term light curve of RX J0838–2827. Fluxes relative to the most recent observations are shown in the inset.

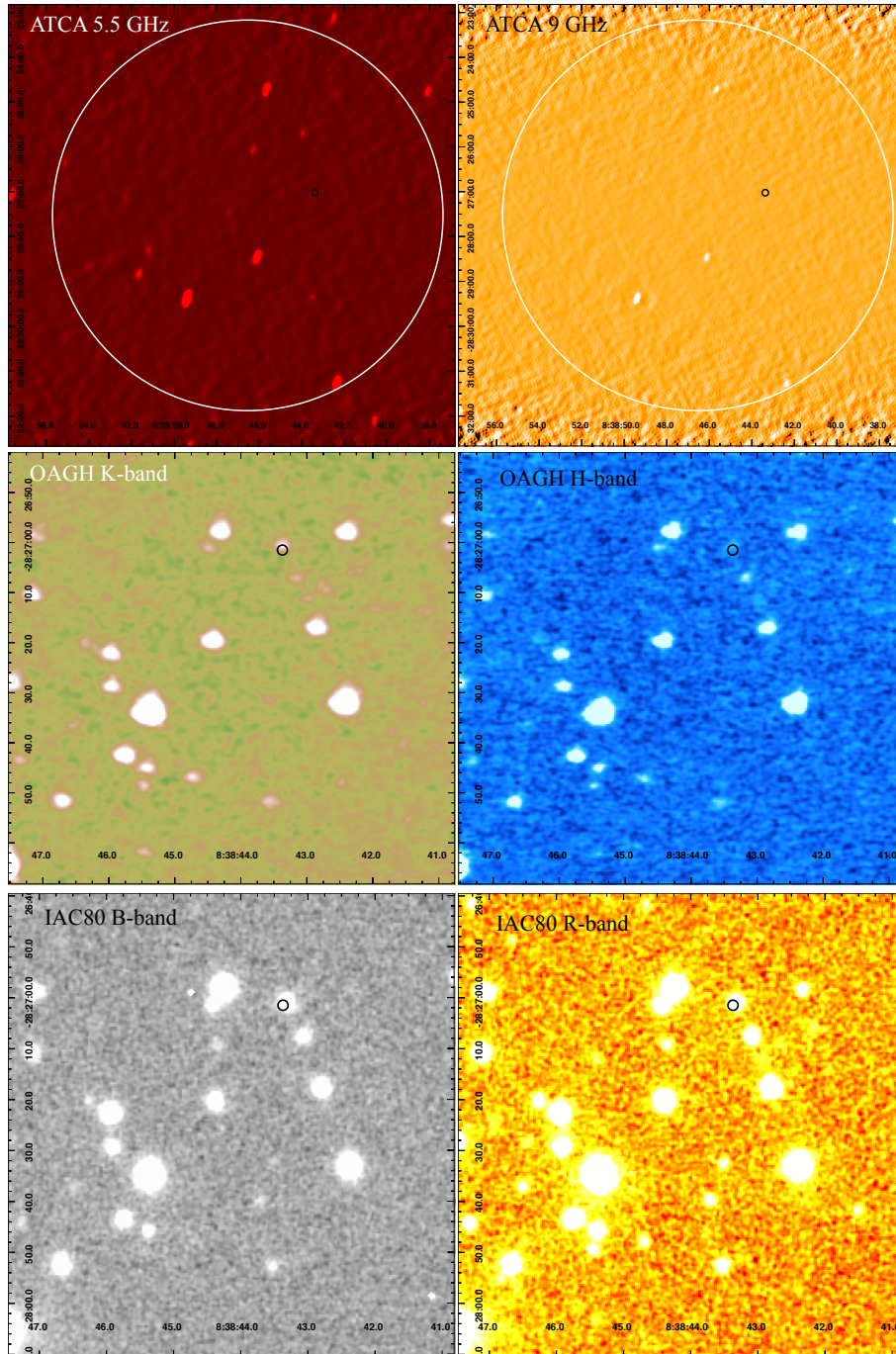


Figure 5. Field around RX J0838–2827 in different energy bands. From top-left to bottom-right we show the radio images taken by ATCA at 5.5 GHz (top left) and 9 GHz (top right) with superimposed the Fermi-LAT 3FGL position (0.033deg error circle at 68% confidence level), and a 2 arcsecond black circle around the optical position of RX J0838–2827. In the middle panels we show the infrared *K* band (middle left) and *H* band (middle right) as observed by the OAGH telescope, and in the bottom panels the optical *B* band (bottom left) and *R* band (bottom right) from the IAC80 telescope. The black circle of the middle and bottom panels is centred on the RX J0838–2827 best optical position with an error radius of 1 arcsec (enlarging by a factor of 5 the optical positional accuracy for imaging purposes). North is up, and east is left.

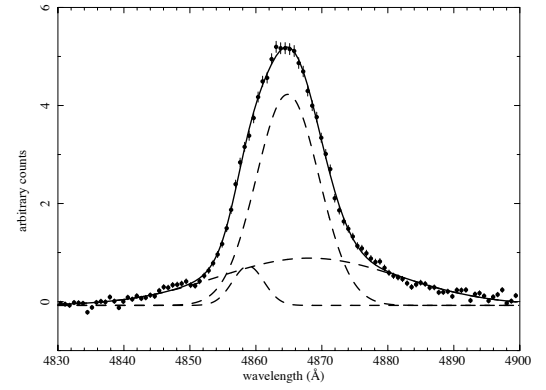
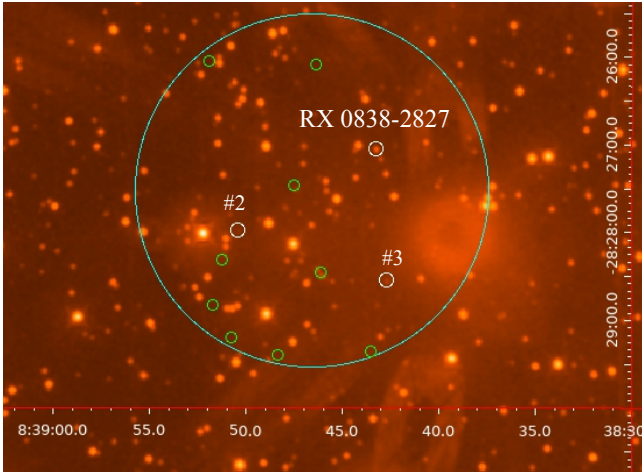
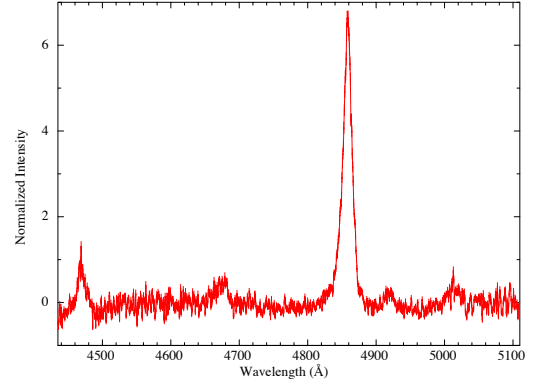
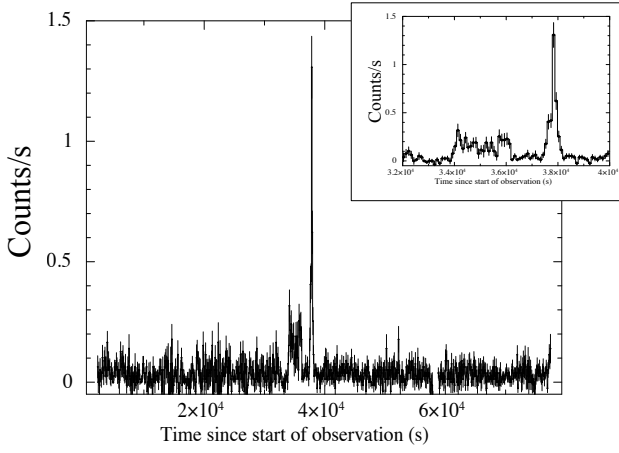


Figure 6. *Top panel:* 0.2–10 keV background-subtracted and exposure-corrected pn light curve of source #2 (XMMJ083850.4–282759) during the second *XMM-Newton* observation. A binning time of 400 s was adopted in the inset with the zoom of the flare. Time is in units of seconds since the start of the exposure. *Bottom panel:* XMM-OM field of view with superimposed the 68% *Fermi*-LAT position of 3FGL J0838.8–2829 (light blue), the three X-ray sources detected in the XMM-pn field of view (white), and in green we report on the position of all the radio detection at 5.5 GHz within the 3FGL error circle (see text for details).

Figure 7. *Top panel:* an example of the ESO-NTT spectra of RX J0838–2827. *Bottom panel:* modelling of the $H\beta$ emission line with three Gaussians (see text for details).

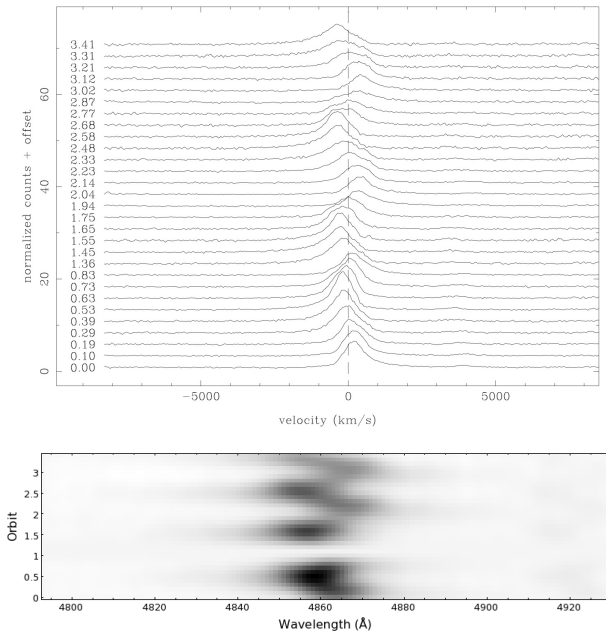


Figure 8. ESO-NTT spectra of RX J0838–2827. Top panel: The H_{β} emission line as a function of the orbital phase (computed assuming phase zero for the first NTT spectrum). Bottom panel: Trailed spectra centered on the H_{β} emission line; orbital phases are relative to the 1.64 hr periodicity.

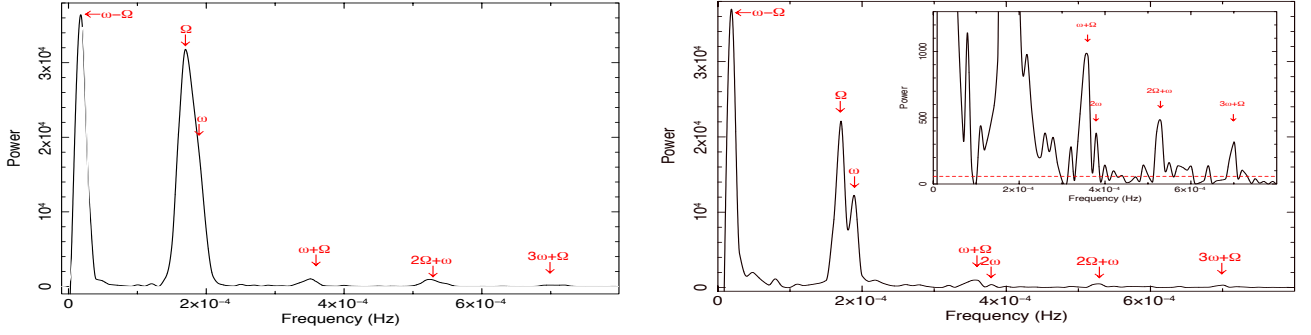


Figure 9. Power spectral distribution of RX J0838–2827 computed from the pn data of the first (left) and the second (right) *XMM-Newton* observations. The frequencies/peaks corresponding to the 1.64h and 1.47 h modulations of the light curves are identified by Ω and ω , respectively, and several other significant peaks can be interpreted as sidebands or higher harmonics of these frequencies (see sect. 5.1). The inset is a zoom of the y-axis for clarity purposes, where the 5σ significance threshold is reported with a dashed line.

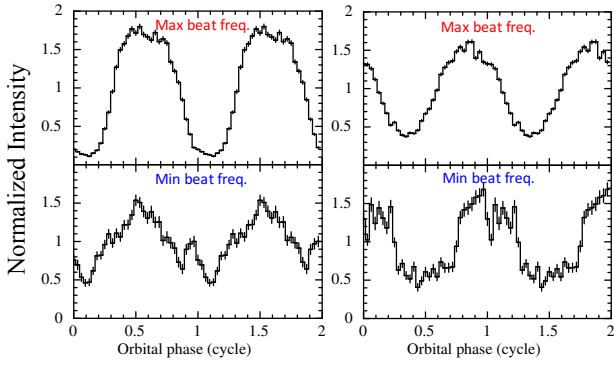


Figure 10. *XMM-Newton* light-curves for the first (left) and second (right) observations folded at the system orbital period. Top panels are relative to the time-span of the maximum of the beat modulation, and bottom panels to the minimum.

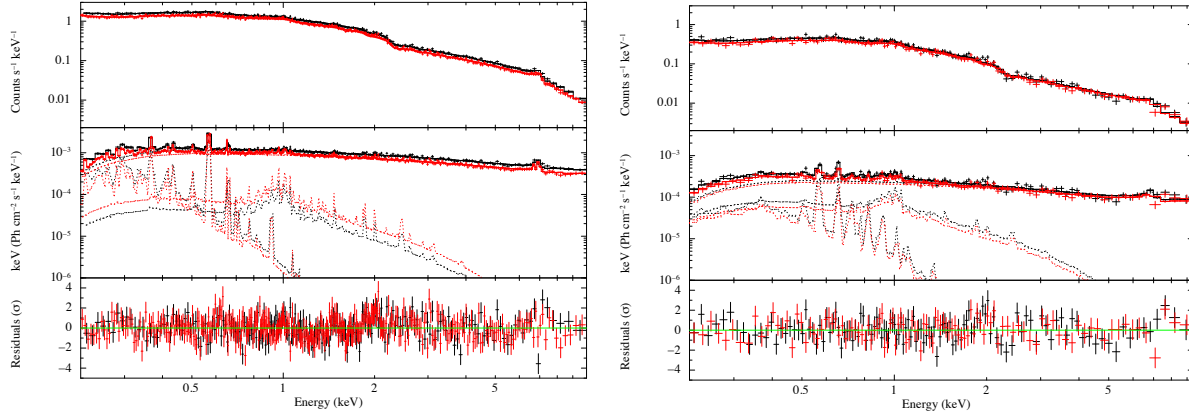


Figure 11. Left panels show the results of the spectral analysis of the peak spectra of the long-term modulation, while on the right we show the results of the minima. From top to bottom (in both panels): 0.2–10 keV pn spectra of RX J0838–2827 fitted to the TBABS*PCFABS*(MEKAL+MEKAL+MEKAL) model (solid lines); $E \times f(E)$ unfolded spectra; post-fit residuals in units of standard deviations. Black color refers to the first observation, red color to the second observation.

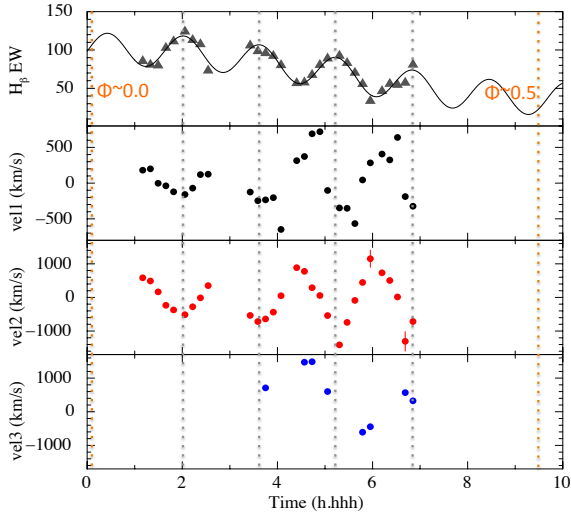


Figure 12. Equivalent width of the $H\beta$ line (top panel), and velocities of three $H\beta$ components as a function of time as observed from ESO-NTT. Vertical dashed-lines report on the phase 0 and 0.5 of the 15.2 h modulation (orange), and on the maxima of the 1.64 h orbital modulation (grey).

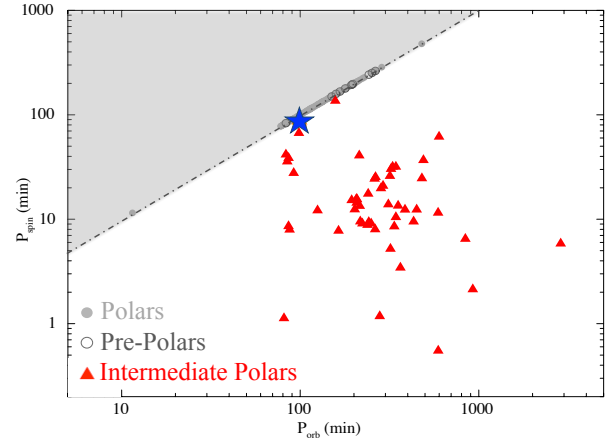


Figure 13. Orbital and spin periods of polars (grey dots), pre-polars (dark grey circles), and IPs (red triangles) known to date, with the values derived for RX J0838–2827 shown as a blue star. The dot-dashed line represent $P_{\text{spin}} = P_{\text{orb}}$. Data adapted from Ferrario, de Martino & Gänsicke (2015).

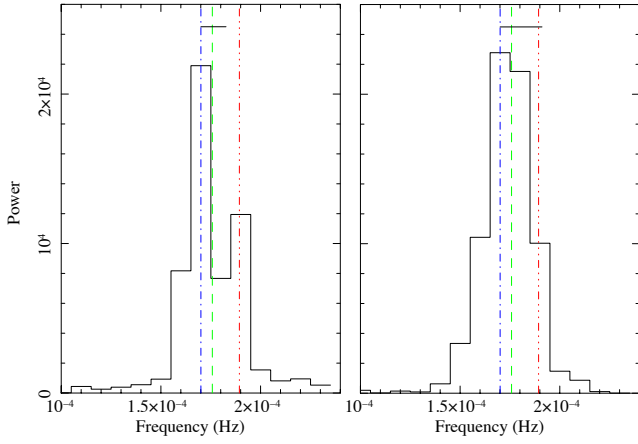


Figure 14. Zoom on the PSD structure around Ω (blue dot-dash line) and ω (red 3dot-dash line) in the second *XMM-Newton* observation (left) and only in the first 45 ks of it (before the flux minimum and the possible phase jump, see Sect. 5.1). The dashed green line indicates the frequency of the signal reported by Halpern et al. (2017). The length of the horizontal bar at the top indicates the Fourier intrinsic resolution of the data segment.